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KIC 10001893: a pulsating sdB star with multiple trapped modes

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ABSTRACT

KIC 10001893 is a V1093 Her type pulsating subdwarf-B (sdB) star, which was observed extensively by the *Kepler* spacecraft. It was a part of the survey searching for compact pulsators in the *Kepler* field of view. An amplitude spectrum of the star demonstrates a rich content of g modes between 102 and 496 μHz as well as a few p modes above 2000 μHz . In total, we found 110 frequencies. The g-mode region contains 104 frequencies, while the p-mode region contains just 6, altogether indicating the hybrid nature of KIC 10001893. The main goal of our analysis was to identify the detected modes and to find some features, which will significantly help modelling the star. We found no multiplets that point at a pole-on orientation; however, we defined modal degrees and relative radial orders using asymptotic period spacing. Eventually, we assigned 32 dipole $l = 1$ and 18 quadrupole $l = 2$ modes. The most remarkable feature we found is trapped modes, which are clearly seen in a reduced-period diagram. It is the second time that three trapped modes are detected in one pulsating sdB star. Since the more trapped modes we find, the better sounding of the stellar interior we can do, this feature provides important constraints on the physical structure of the star.

Key words: stars: oscillations – subdwarfs.

1 INTRODUCTION

Subdwarf-B (sdB) stars are located between the horizontal branch (HB) and the white-dwarf cooling track, on the so-called extreme horizontal branch (EHB, Heber 2016). The sdB stars are compact objects with masses typically around 0.5 solar masses and have surface gravities between $\log g$ of 5.0 and 5.8, which correspond to radii between 0.15 and 0.35 R_{\odot} . Observed effective temperatures range from about 20 000 to 40 000 K. These stars are core He-burning with thin hydrogen envelopes ($M_{\text{env}} < 0.01 M_{\odot}$). Such a small mass of the hydrogen envelope does not allow sdB stars to ascend the asymptotic giant branch. Therefore, after all helium is exhausted in their cores, sdB stars move directly to the white dwarf cooling track. The most enigmatic part of this evolution is the envelope-stripping phase, during which almost the entire hydrogen envelope is removed. This must occur close to the tip of the red giant branch, just before the occurrence of the helium flash, after which the envelope contracts while the core expands, and before the star settles down in its stable helium-burning phase. The cause of the envelope stripping is most likely mass transfer involving a close companion to an sdB star (Han et al. 2002), or mass-loss

from a significantly enhanced stellar wind (Yong, Demarque & Yi 2000).

The first pulsating sdB was discovered by a South African group of astronomers (Kilkenny et al. 1997) and enables us to probe the interiors of sdB stars using asteroseismology. It is a powerful tool that allows us to sound stellar structure using natural vibrations. Pulsating sdB (sdBV) stars are classified as V361 Hya or V1093 Her stars based on their periods. The former class is dominated by p-mode frequencies higher than 2000 μHz while the latter class is dominated by g modes, which are usually below 1000 μHz . Some sdBV stars show both kinds of modes and are therefore called hybrid stars. All hybrids detected from the ground were p mode dominated (Baran et al. 2005; Schuh et al. 2006), while those observed from space are mostly g mode dominated (KIC 10001893 being an example), with a few exceptions of p-mode-dominated ones (e.g. Baran et al. 2012). The first theoretical models of sdBV stars were a paper by Charpinet et al. (1996) followed by, for example, Charpinet et al. (1997, 2000) and Fontaine et al. (2003). A successful application of those models can be found in a number of papers (e.g. Charpinet et al. 2011; Van Grootel et al. 2013).

During the last six years significant headway has been made in the field of sdBV stars. The *Kepler* spacecraft delivered unprecedented data revealing features that were rarely or never seen in ground-based data. These features include rotational multiplets

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(e.g. Baran 2012), asymptotic period spacing (e.g. Reed et al. 2011), Doppler beaming (e.g. Telting et al. 2012) and mode trapping (Østensen et al. 2014; Foster et al. 2015; Kern et al. 2017).

The *Kepler* observations made it possible to resolve the g-mode region in the V1093 Her stars, something that, so far, has not been possible with ground-based data. This had profound impact on the theoretical models, where the physics related to mixing (especially diffusion and overshooting) is poorly constrained. The first results by Reed et al. (2011) indicated that the asymptotic sequences were much smoother than predicted by the early theoretical works, in which stratification played a strong role and produced a sharp boundary in the H/He transition zone. Such sharp boundaries are very efficient in producing trapped modes, and it was speculated that the absence of trapped modes might indicate that stronger diffusion could reduce the impact of this boundary. However, Charpinet et al. (2014) demonstrated that while the boundary indeed produces significant trapping at low radial order, modes of higher order (which are the ones that appear with high amplitude in the observations) become less sensitive to this transition layer as their wavelengths become long compared to the size of the boundary. But this discussion was abruptly turned on its head with the first discovery of trapped modes in KIC 10553698A by Østensen et al. (2014). With such a clear signature of mode trapping, it became possible for the modellers to test the adequacy of their models in reproducing the observed features. The first result was reported already by Constantino et al. (2015), where they explored models with different degrees of core overshooting parameters and found that their models could reproduce similar trapping effects caused by the sharp composition gradient at the edge of the partially mixed zone associated with the C-O/He boundary outside the convective inner core. More recently, Ghasemi et al. (2017) explored models with various combinations of overshooting and diffusion parameters with the particular aim of matching the trapping patterns seen in KIC 10553698A. In their scenarios with small and moderate overshooting, the mixing leads to the emergence of convective shells around the core, which produces mode-trapping patterns that have comparable trapping signatures to those observed in KIC 10553698A.

In this paper, we present our analysis of KIC 10001893, which was observed with the *Kepler* spacecraft. It was confirmed as a V1093 Her pulsator with one month of data (Q3.3) during the *Kepler* survey phase (Østensen et al. 2010, 2011; Baran et al. 2011). Baran et al. (2011) found 27 frequencies, mostly in the g-mode region, with asymptotic sequences as reported by Reed et al. (2011). Silvotti et al. (2014) used data covering Q6–Q17.2 (1051 d), which is the entire (continuous) data coverage available. However, they concentrated only on the lowest frequency region, below the so-called *cut-off* frequency and interpreted three frequencies, found in that region, as indications of exoplanets.

Here, we have used the same 3-yr data set as Silvotti et al. (2014) (Section 2), and we provide a detailed analysis of the frequencies above the cut-off frequency. The increased span of the observations gives better resolution and lowers the detection limit compared to Baran et al. (2011), allowing us to detect more frequencies (Section 3). In Sections 4–6, we apply seismic tools to the pulsations to identify features useful for constraining models, in particular the échelle and reduced period-diagrams. We provide our conclusions in Section 7.

2 PHOTOMETRIC DATA

KIC 10001893 was observed by the *Kepler* spacecraft. The observations started on 2010 June 24 and finished on 2013 May 11, which

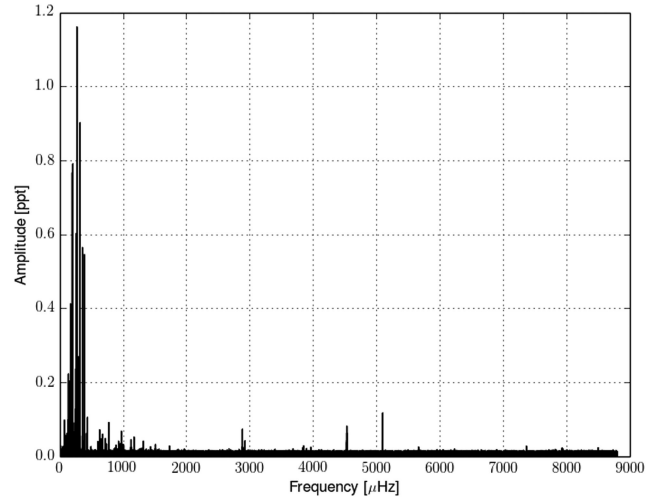


Figure 1. Amplitude spectrum, calculated up to the Nyquist frequency, showing both g and p modes. Close-ups of both regions are shown in Figs 2 and 3.

covered Q6 to Q17.2. We downloaded all available data from the ‘Barbara A. Mikulski Archive for Space Telescopes’ (MAST).¹ The *Kepler* spacecraft has two types of exposure times, which are short cadence (~ 1 min) and long cadence (~ 30 min). We used short-cadence data that cover the frequency range up to the Nyquist at 8495 μHz , and assuring both p- and g-mode regions are covered. The MAST data base provides data processed in a couple of different ways (RAW, SAP, PDCSAP). We used fluxes with the pre-search data conditioning (PDCSAP) module, which are widely known as PDC fluxes. These fluxes are corrected for systematic errors and contamination from nearby stars, while the other fluxes are not. We clipped data at 4σ and de-trended each monthly chunk of data separately. Finally, we stitched the monthly chunks together making the data ready for Fourier analysis.

3 FOURIER ANALYSIS

We used a Fourier technique to calculate the amplitude spectrum from Q6–17.2 data and we show it in Fig. 1. The frequency resolution equals 0.0162 μHz as defined by $1.5/T$, where T is the time coverage of the data (Baran 2012). We first tried the pre-whitening technique to fit and remove frequencies; however, profiles of signals at many frequencies were complex and we were unable to fit those. Therefore, we decided to determine the frequencies by-eye. We applied a 5.5σ detection threshold (Baran, Koen & Pokrzywka 2015a) of 0.022 ppt (parts per thousand). We checked the detected frequencies against the list of artefacts (Baran 2013). Eventually, we found 110 significant peaks above the detection threshold, showing KIC 10001893 to be a rich g-mode pulsator with most pulsations below 500 μHz but significant frequencies all the way up to 3897 μHz . We list all detected frequencies in Table 1.

3.1 p modes

We found six frequencies in the high-frequency region. This area is relatively sparse in frequencies, with all of them located between 2880 and 3900 μHz . We show a close-up of this region in Fig. 2. The highest amplitude is 0.07 ppt. We assigned those frequencies

¹ archive.stsci.edu

Table 1. A list of frequencies found in the amplitude spectrum of KIC 10001893. Trapped modes are indicated with a ‘t’ letter. Asterisks mark amplitudes between 4.75 and 5.5σ .

ID	Frequencies (μHz)	Period (s)	Amplitude (ppt)	<i>l</i>	<i>n</i>
<i>f</i> ₁	76.456 10	13 079.4	0.026		
<i>f</i> ₂	77.529 59	12 898.3	0.098		
<i>f</i> ₃	78.455 98	12 746.0	0.029		
<i>f</i> ₄	79.716 21	12 544.5	0.023		
<i>f</i> ₅	81.979 98	12 198.1	0.040		
<i>f</i> ₆	83.938 38	11 913.5	0.021*		
<i>f</i> ₇	87.361 42	11 446.7	0.024		
<i>f</i> ₈	89.088 44	11 224.8	0.048		
<i>f</i> ₉	91.749 85	10 899.2	0.024		
<i>f</i> ₁₀	96.938 67	10 315.8	0.056		
<i>f</i> ₁₁	98.980 50	10 103.0	0.020*		
<i>f</i> ₁₂	102.172 18	9787.4	0.027	1	36
<i>f</i> ₁₃	105.019 95	9522.0	0.031	1	35
<i>f</i> ₁₄	107.833 50	9269.3	0.031		
<i>f</i> ₁₅	108.189 98	9243.0	0.036	1	34
<i>f</i> ₁₆	111.644 52	8957.0	0.023		
<i>f</i> ₁₇	114.416 47	8740.0	0.062		
<i>f</i> ₁₈	114.710 47	8717.6	0.038	1	32
<i>f</i> ₁₉	116.912 57	8553.4	0.060		
<i>f</i> ₂₀	118.333 39	8450.7	0.032	1	31
<i>f</i> ₂₁	120.494 99	8299.1	0.022		
<i>f</i> ₂₂	122.076 27	8191.6	0.019*	1	30
<i>f</i> ₂₃	123.889 63	8071.7	0.023		
<i>f</i> ₂₄	125.855 81	7945.6	0.048	1	29
<i>f</i> ₂₅	126.890 67	7880.8	0.034	1	t
<i>f</i> ₂₆	130.550 00	7659.9	0.041	1	28
<i>f</i> ₂₇	132.474 89	7548.6	0.033		
<i>f</i> ₂₈	135.328 98	7389.4	0.098	1	27
<i>f</i> ₂₉	140.520 48	7116.4	0.223	1	26
<i>f</i> ₃₀	141.474 73	7068.4	0.029		
<i>f</i> ₃₁	145.736 47	6861.7	0.036		
<i>f</i> ₃₂	146.100 57	6844.6	0.088	1	25
<i>f</i> ₃₃	152.061 18	6576.3	0.098	1	24
<i>f</i> ₃₄	155.967 30	6411.6	0.028		
<i>f</i> ₃₅	158.082 77	6325.8	0.204	1	23
<i>f</i> ₃₆	160.058 90	6247.7	0.048	1	t
<i>f</i> ₃₇	163.054 99	6132.9	0.033		
<i>f</i> ₃₈	164.991 99	6060.9	0.069		
<i>f</i> ₃₉	165.395 87	6046.1	0.123	1	22
<i>f</i> ₄₀	166.212 35	6016.4	0.045		
<i>f</i> ₄₁	172.458 39	5798.5	0.063		
<i>f</i> ₄₂	172.720 52	5789.7	0.413	1	21
<i>f</i> ₄₃	173.190 16	5774.0	0.057		
<i>f</i> ₄₄	176.678 44	5660.0	0.034		
<i>f</i> ₄₅	178.794 92	5593.0	0.021*		
<i>f</i> ₄₆	180.554 30	5538.5	0.065	1	20
<i>f</i> ₄₇	181.943 88	5496.2	0.039	2	35
<i>f</i> ₄₈	187.265 91	5340.0	0.019*	2	34
<i>f</i> ₄₉	189.093 10	5288.4	0.076	1	19
<i>f</i> ₅₀	192.715 35	5189.0	0.033	2	33
<i>f</i> ₅₁	196.417 34	5091.2	0.036		
<i>f</i> ₅₂	199.111 96	5022.3	0.770	1	18
<i>f</i> ₅₃	201.174 86	4970.8	0.034	1	t
<i>f</i> ₅₄	204.695 71	4885.3	0.097	2	31
<i>f</i> ₅₅	210.681 55	4746.5	0.793	1	17
<i>f</i> ₅₆	211.188 78	4735.1	0.040	2	30
<i>f</i> ₅₇	217.755 80	4592.3	0.025	2	29
<i>f</i> ₅₈	219.408 91	4557.7	0.046	2	t
<i>f</i> ₅₉	224.719 10	4450.0	0.054	1	16
<i>f</i> ₆₀	225.896 81	4426.8	0.027	2	28
<i>f</i> ₆₁	234.175 58	4270.3	0.066	2	27
<i>f</i> ₆₂	238.982 88	4184.4	0.091	1	15

Table 1 – *continued*

ID	Frequencies (μHz)	Period (s)	Amplitude (ppt)	<i>l</i>	<i>n</i>
<i>f</i> ₆₃	243.143 35	4112.8	0.052	2	26
<i>f</i> ₆₄	252.793 36	3955.8	0.129	2	25
<i>f</i> ₆₅	255.167 13	3919.0	0.236	1	14
<i>f</i> ₆₆	262.977 96	3802.6	0.602	2	24
<i>f</i> ₆₇	273.336 06	3658.5	0.086	2	23
<i>f</i> ₆₈	274.303 26	3645.6	1.162	1	13
<i>f</i> ₆₉	276.357 60	3618.5	0.053	2	t
<i>f</i> ₇₀	286.000 28	3496.5	0.186	2	22
<i>f</i> ₇₁	298.596 59	3349.0	0.051	2	21
<i>f</i> ₇₂	298.667 94	3348.2	0.270	1	12
<i>f</i> ₇₃	312.061 16	3204.5	0.039	2	20
<i>f</i> ₇₄	323.981 07	3086.6	0.902	1	11
<i>f</i> ₇₅	326.786 70	3060.1	0.053	2	19
<i>f</i> ₇₆	359.686 35	2780.2	0.565	1	10
<i>f</i> ₇₇	391.374 11	2555.1	0.546	1	9
<i>f</i> ₇₈	426.039 53	2347.2	0.062		
<i>f</i> ₇₉	440.509 22	2270.1	0.105	1	8
<i>f</i> ₈₀	496.154 80	2015.5	0.026	1	7
<i>f</i> ₈₁	605.363 52	1651.9	0.040		
<i>f</i> ₈₂	633.753 72	1577.9	0.072		
<i>f</i> ₈₃	660.894 85	1513.1	0.048		
<i>f</i> ₈₄	679.024 92	1472.7	0.060		
<i>f</i> ₈₅	721.032 51	1386.9	0.049		
<i>f</i> ₈₆	742.059 95	1347.6	0.034		
<i>f</i> ₈₇	777.846 91	1285.6	0.091		
<i>f</i> ₈₈	883.938 83	1131.3	0.022		
<i>f</i> ₈₉	892.299 45	1120.7	0.030		
<i>f</i> ₉₀	924.727 20	1081.4	0.021*		
<i>f</i> ₉₁	930.925 33	1074.2	0.041		
<i>f</i> ₉₂	943.663 30	1059.7	0.026		
<i>f</i> ₉₃	966.090 23	1035.1	0.035		
<i>f</i> ₉₄	977.517 10	1023.0	0.068		
<i>f</i> ₉₅	1004.520 34	995.5	0.032		
<i>f</i> ₉₆	1013.171 22	987.0	0.030		
<i>f</i> ₉₇	1128.413 45	886.2	0.045		
<i>f</i> ₉₈	1175.778 95	850.5	0.052		
<i>f</i> ₉₉	1267.587 78	788.9	0.021*		
<i>f</i> ₁₀₀	1294.665 97	772.4	0.022		
<i>f</i> ₁₀₁	1321.877 06	756.5	0.041		
<i>f</i> ₁₀₂	1435.750 17	696.5	0.026		
<i>f</i> ₁₀₃	1513.317 19	660.8	0.032		
<i>f</i> ₁₀₄	1734.304 54	576.6	0.028		
<i>f</i> ₁₀₅	2884.338 04	346.7	0.073		
<i>f</i> ₁₀₆	2898.550 72	345.0	0.023		
<i>f</i> ₁₀₇	2925.687 53	341.8	0.042		
<i>f</i> ₁₀₈	3834.355 82	260.8	0.023		
<i>f</i> ₁₀₉	3849.114 70	259.8	0.028		
<i>f</i> ₁₁₀	3897.116 13	256.6	0.021*		

with p modes. No multiplets are present in this region, and since such features are one of the most useful tools to make mode identifications, we were left with no chance to directly identify individual modes. The multiplets can also be useful in estimating stellar rotation. An absence of split modes can happen if the stellar rotation axis is pointed towards the observer, which efficiently suppresses the amplitudes of the side components, or if the rotation rate is so slow that the span of the data is insufficient to resolve the multiplets. More details will be given in Section 4. A few frequencies beyond $4000 \mu\text{Hz}$ are of instrumental origin, related to the widely known long cadence-readout time.

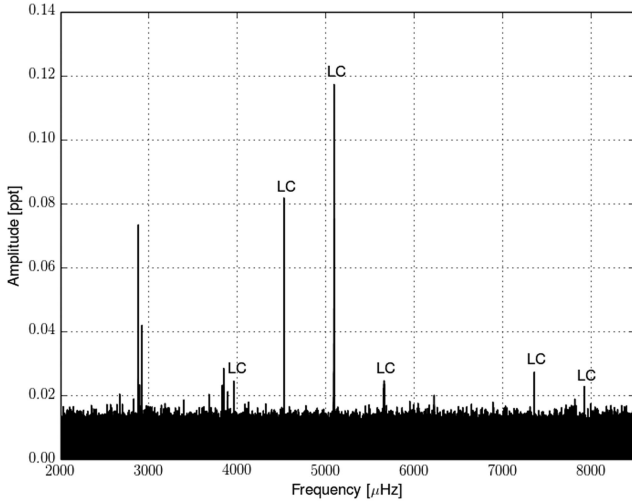


Figure 2. A close-up of the amplitude spectrum showing the high-frequency region. LC stands for long-cadence-readout-time artefacts.

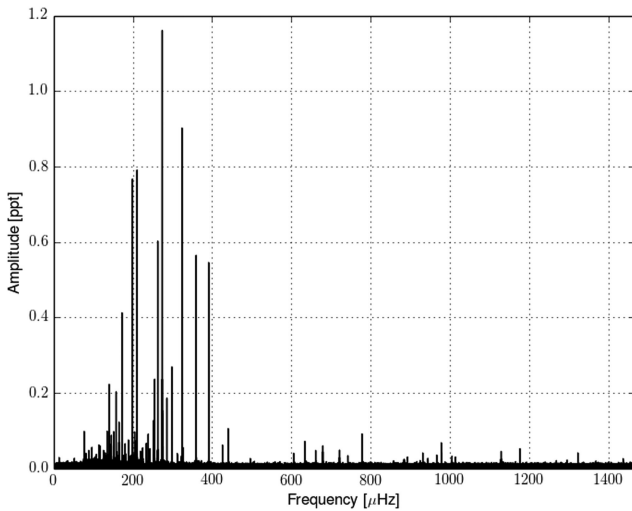


Figure 3. A close-up of the amplitude spectrum showing the low-frequency region.

3.2 g modes

We show the g-mode region of the amplitude spectrum in Fig. 3. This region is rich in frequencies with relatively large amplitudes compared to the p modes (see Fig. 1). The frequencies with the highest amplitudes are located below 400 μHz . The region above 500 μHz up to 2000 μHz shows numerous low-amplitude peaks that do not follow the spacing sequences of the low-order modes. Although modes are not expected to follow the asymptotic spacing at low n , there seem to be too many of them for all to be of degree $l = 1$ or 2. In several other *Kepler*-observed sdBV stars (Telting et al. 2014; Foster et al. 2015; Kern et al. 2017), multiplets in this frequency range indicate $3 \leq l \leq 9$ modes. A few frequencies are of instrumental nature (566.43 μHz and its multiples). In total, we found 24 frequencies in this region. But since we cannot see any signs of rotational splitting, there is no chance of identifying these modes.

4 MULTIPLETS

The ultimate goal of our analysis was to identify modes to constrain theoretical models of this star. Mode identifications describe a mode’s geometry, i.e. a radial order n , a modal degree l and an azimuthal order m of the mode. Information on these parameters is very crucial for calculation of stellar interior. Each mode is described by three numbers and therefore each mode adds three unknown parameters to the model. Since the more the free parameters there are, the less reliable the model is, it is important to identify as many modes as possible.

In the presence of stellar rotation, non-radial modes of degree l split into $2l + 1$ components differing in m number. In the case of slow rotation, i.e. rotational period is much higher than a pulsation period, the azimuthal frequencies can be derived from the following equation:

$$\nu_{n,l,m} = \nu_{n,l,0} + \Delta\nu_{n,l,m} = \nu_{n,l,0} + m \frac{1 - C_{n,l}}{P_{\text{rot}}}, \quad (1)$$

where $\Delta\nu_{n,l,m}$ is a rotational splitting. P_{rot} is a star’s rotation period and $C_{n,l}$ is the Ledoux constant. We searched for multiplets among g modes with a null result. The most likely explanation for this is that the inclination is too low, as was inferred by Silvotti et al. (2014). The authors concluded that the inclination may be as low as just a few degrees. If so, KIC 10001893 is another example of a pole-on oriented sdB star (such as KIC 8302197; Baran et al. 2015b). Other possibilities, like extremely slow rotation or selective driving of $m = 0$ modes cannot be completely ruled out, though we consider this to be unlikely. The lack of multiplets prevents us from identifying azimuthal orders and determining a rotation rate for KIC 10001893.

If the absence of rotational splitting is due to extremely slow rotation, the time resolution of the Kepler data set provides a limit on the rotation period of $P_{\text{rot}} \geq 715$ d. Here, we used the full width at half-maximum of a few amplitude-stable modes (0.00116 c/d) to approximate the resolution needed to tell apart an $l > 1$ multiplet, and we assumed the δm splitting of $(1 - C_{nl})/P_{\text{rot}} \geq 0.83/P_{\text{rot}}$, expected for high order $l > 1$ g modes. We examined the amplitudes of many of our modes and do not see a single characteristic time-scale of amplitude variability longer than 715 d, which one may expect for unresolved beating between multiplet components. Therefore, if the lack of multiplets is due to very slow rotation, the rotation period must be well longer than the length of the data set. This is in contrast with typical rotational rates of tens of days detected thus far (e.g. Baran 2012; Baran & Winans 2012; Telting et al. 2012; Østensen et al. 2014; Reed et al. 2014; Foster et al. 2015).

For the other known case of an sdB pulsator that does not show any multiplets in a 3-year data set, KIC 8302197, Baran et al. (2015b) argued that the lack of multiplets can be due to geometrical cancellation if the inclination is lower than a few degrees, but that the probability of having such low-inclination angles is such that one would expect only a few such cases in every 1000 sdB stars. Finding two such cases in a sample of only ~ 16 well-studied sdB pulsators, is therefore somewhat problematic.

5 PERIOD SPACING

In the asymptotic limit $n \gg l$, consecutive radial overtones are evenly spaced in period (e.g. Reed et al. 2011). For given n and l values, periods of consecutive overtones can be calculated from

$$P_{l,n} = \frac{\Delta\Pi}{\sqrt{l(l+1)}}n + \epsilon_l, \quad (2)$$

where $\Delta\Pi$ is the reduced-period spacing (see below) and ϵ_l is a constant offset smaller than $\Delta\Pi$. With this equation, we define the asymptotic radial order of a mode, n , but note that this does not allow us to derive the true radial order of a mode, which is usually denoted k , since any trapped modes will be inserted into the mode sequence. Thus, each trapped mode will bump the mode order k by one relative to the asymptotic order n , at the point in the sequence where it occurs. The reduced-period spacing $\Delta\Pi$ can be inferred from the échelle diagrams by converting the observed period spacings for each l , ΔP_l , to reduced period, $\Delta\Pi = \sqrt{l(l+1)}\Delta P_l$. The offset ϵ_l can also be derived from the échelle diagrams as the mean value of the observed periods associated with a given l modulo the period spacing. It was found from observations that the period spacing of $l = 1$ modes is around 250 s (Reed et al. 2011), matching predictions from theoretical models (Charpinet et al. 2000). The spacings of $l = 2$ modes are found to be around 145 s, so that both produces reduced spacings of ~ 350 s. Period spacings for $l = 3$ would translate to around 102 s, but has never been detected.

Since pulsation periods depend on the local sound speed, chemical stratification in the resonant cavity will cause fluctuations in the observed period spacings at low n , but these should disappear at higher n (Charpinet et al. 2014). For KIC 10001893, we derived the average of the spacing value for $l = 1$ to be 268.0 ± 0.5 and 153 ± 0.4 s, respectively, for $l = 1$ and 2. Our $l = 1$ and 2 values match those of Reed et al. (2011). The sequence of $l = 1$ modes is fairly complete, with 32 overtones identified. While 18 $l = 2$ overtones were identified, the sequence is nearly complete, only missing $n = 32$. The identification of modal degrees is marked in Fig. 4 and provided in Table 1. KIC 10001893 is therefore another very successful case for mode identification in sdB stars, providing significant constraints for stellar modelling. Although the multiplets are not present, their absence actually makes it simpler since we can associate all modes with $m = 0$ and avoid the ambiguity that exists when having incomplete multiplets, as was seen in KIC 10553698A (Østensen et al. 2014). A pole on orientation means that only the central components are visible, while an extremely slow rotation would imply that modes of different m have the same frequency, effectively producing the same result. In either case, this does not reduce the value of the identification that we are able to make from asymptotic sequences alone. Although, we would have liked to use the multiplets for confirmation of our mode designations, having obtained those verifications from a large number of cases already, we are confident that the asymptotic sequences alone are sufficient to derive reliable mode identifications.

Reed et al. (2011) found period spacings between 230 and 270 s for the g-mode pulsators in the *Kepler* field. The spacing is determined by the size of the pulsation resonant cavity and is the first asteroseismic diagnostic that can be used to constrain pulsation models. For instance Schindler, Green & Arnett (2015) found that models computed with MESA produced much larger cores when type II opacities were included and either overshoot or diffusion was turned on. Still, such large cores would reduce the size of the propagation cavity for g modes and these are incompatible with all but the shortest period spacings found by Reed et al. (2011). Constantino et al. (2015) also concluded that their models tended to produce period spacings significantly below the average inferred from observations. The period spacing is therefore an important constraint for theoretical models, and since KIC 10001893 has one of the largest period spacings for known g-mode pulsators, it is an important case for modelling the extent of resonant cavities.

6 MODE TRAPPING

The asymptotic relation for g modes, equation (2), holds for idealized stellar models with homogeneous composition. In that case the boundaries for a g mode being a standing wave is a surface and, usually, a convective core. The consecutive overtones are perfectly spaced making an even sequence of modes of a given modal degree l , and the échelle diagrams would show a vertical ridge. However, real sdB stars are compositionally stratified. Composition discontinuities will produce steep gradients in density, which may act as a boundary for stellar pulsations. In the sdB stars, the two important composition transition regions are the H/He boundary and the deep transition region between the helium mantle and the central convective core. In the models of Charpinet et al. (2014) it is the H/He transition that produces the strongest trapping effect, while in the work of Ghasemi et al. (2017) strong trapping effects are produced by convective shells forming in the transition zone between the C/O-enhanced convective core and the radiative He mantle. These boundaries can cause partial reflection of the waves, or even act as an extra reflection surface that causes extra modes to be inserted into the otherwise roughly evenly spaced sequence of frequencies.

While the detection of trapped modes is in itself interesting, finding several of them and thereby being able to estimate the spacing between consecutive trapped modes is even more useful since the spacing between consecutive trapped modes can be a particularly important diagnostic when comparing observations with theoretical models. When searching for trapped modes, it is convenient to start with the identifications made in the échelle diagram for the $l = 1$ and 2 sequences and plot them in a reduced-period diagram (where period, P , is converted to reduced period $\Pi = P\sqrt{\ell(\ell+1)}$). After conversion to reduced period, modes of the same order should fall roughly on top of each other. In the échelle diagram, trapped modes are shifted off the ridge and cannot therefore be assigned to a particular l . But with some trial and error it is possible to use the fact that the reduced-period differences should be the same for different l to find solutions where two observed modes correspond to the same order n for different l and therefore line up in the reduced-period diagram.

We calculated échelle diagrams for $l = 1$ up to 5, though only $l = 1$ and 2 overtones show evidence of systematic ridges (shown in Fig. 5). The diagram of $l = 1$ modes shows a ‘hook’ feature observed in some other sdBV stars (e.g. Baran & Winans 2012). Such a feature appears in some models (e.g. Fig. 1 of Charpinet et al. 2014), and is related to the chemical stratification, but a quantitative explanation of this feature has not been presented; therefore, we will postpone a discussion of this feature until models for this star are available. Since there are significant deviations from the ridge towards higher periods (the ‘hook’ feature), we expect the true period spacings to be slightly higher than the means. Therefore, for $l = 1$, we choose a period spacing of 267.5 s, which produces a reasonably vertical ridge when the most outlying points are ignored. For $l = 2$, we enforce the theoretically predicted relationship between $l = 1$ and 2, which gives a period spacing of $267.5/\sqrt{3} = 154.4$ s. We used the asymptotic relation to define n values for our modes, and these values are listed in Table 1. We note that there may be a zero-point offset n_l between the listed n values and the actual mode order k , an offset that we cannot determine without the appropriate modelling. The trapped modes are effectively inserted between consecutive asymptotic radial orders, implying smaller period spacings between adjacent radial orders around the trapped mode.

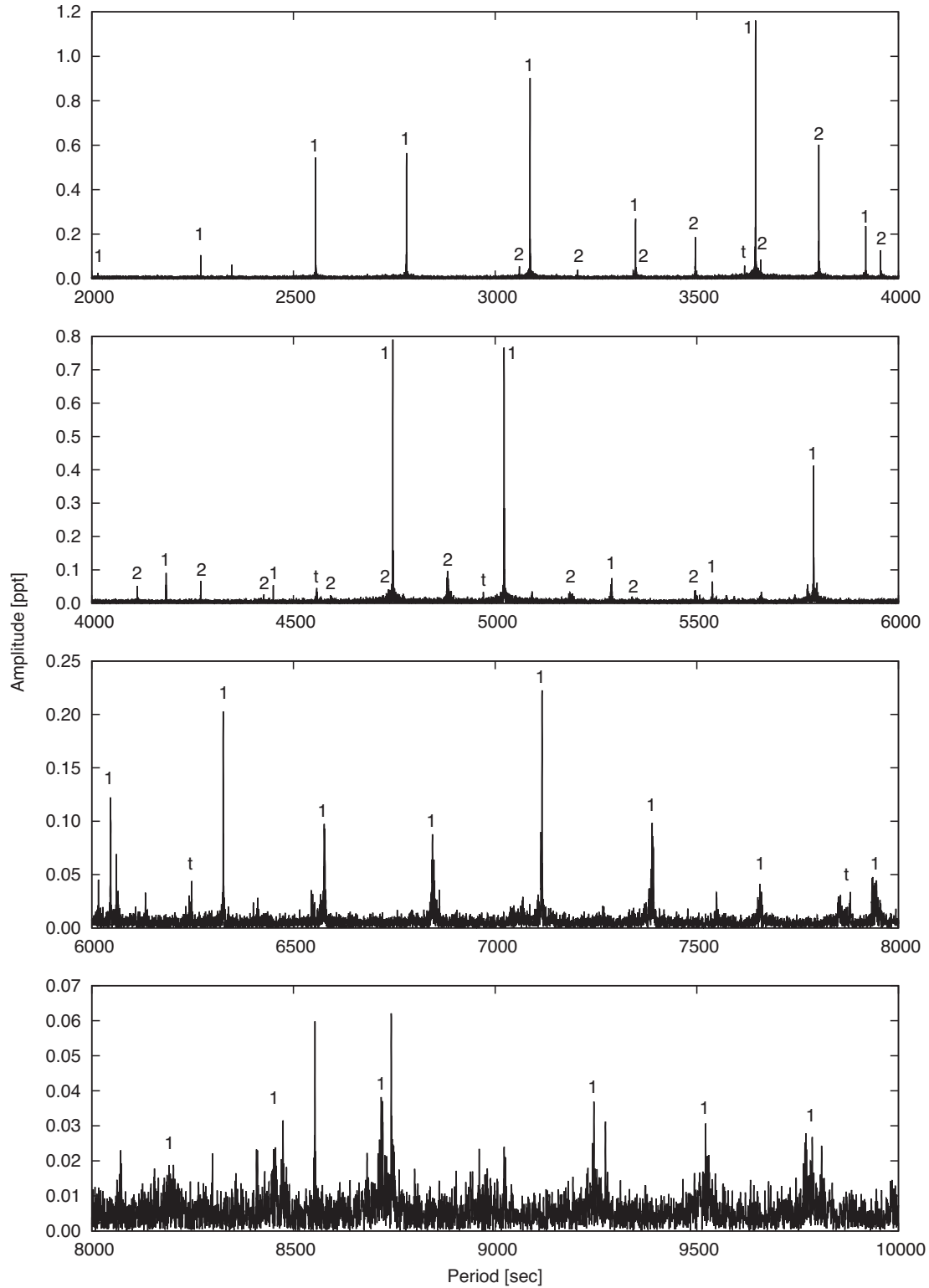


Figure 4. Modal degree assignment. Note that the horizontal axis is in period instead of frequency.

The identification of two trapped modes between radial orders $n = 22$ and 23 , and 28 and 29 , is very convincing as they are well matched in both the $l = 1$ and 2 sequences (Fig. 6). The radial orders around the trapped modes agree between both modal degrees, as predicted by Charpinet et al. (2000) (their equation 30). A possible third trapped mode around $n = 18$ is present only in

the $l = 1$ sequence, and as we could not recover the period-spacing relation for $l = 2$ in that region, the significance of that trapped mode is less convincing. Indeed, there are other unidentified modes that could be associated with trapped modes, but without multiplets, we have no way of preferring one over the other. For the overlap region, however, the match is exceptionally good, which

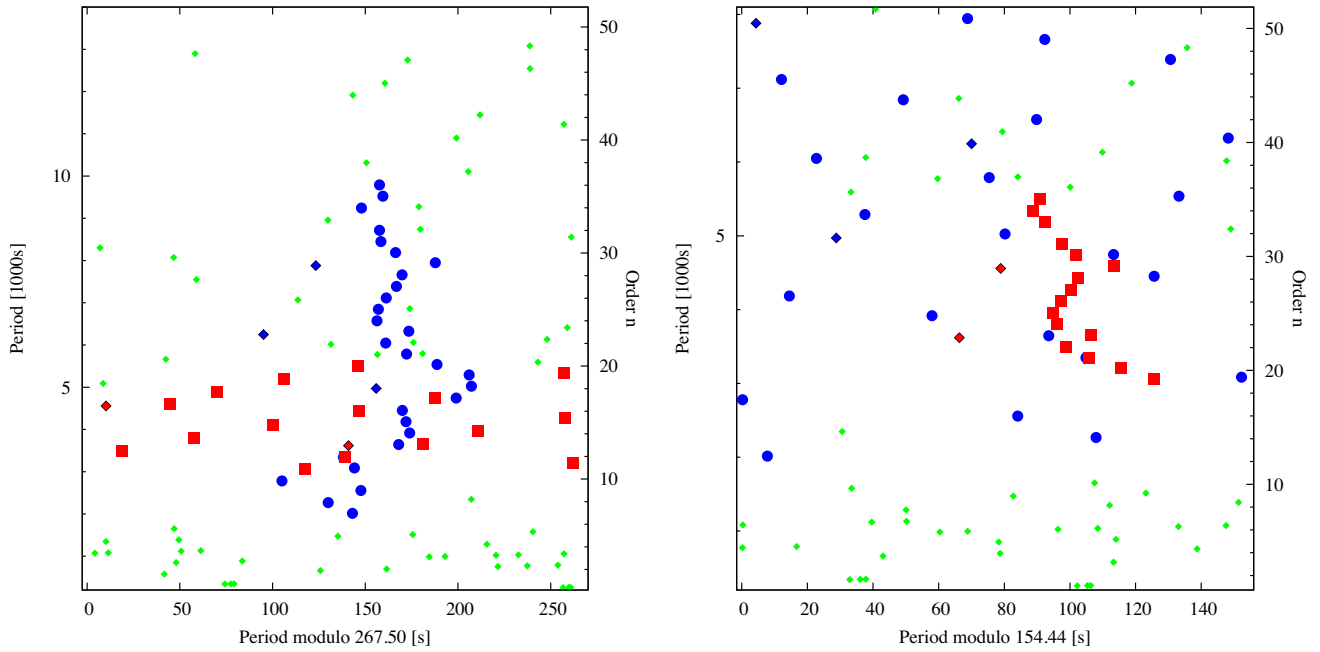


Figure 5. Échelle diagrams folded over the average $l = 1$ (left-hand panel) and $l = 2$ (right-hand panel) period spacing (see text for explanation). Blue dots represent dipole modes, red squares quadrupole modes, while diamonds with black outlines represent trapped modes. Green diamonds show unidentified modes.

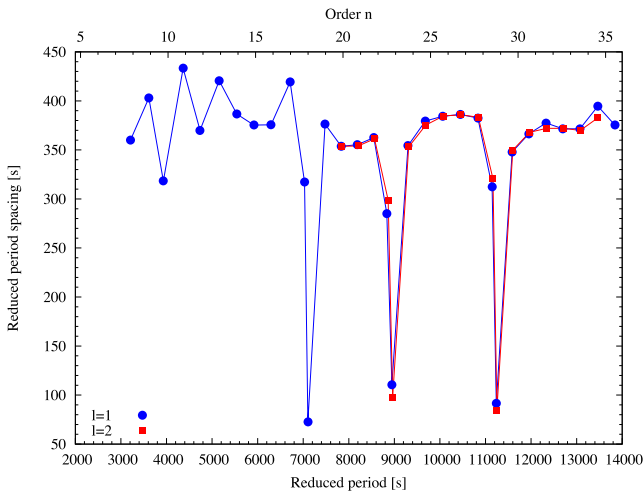


Figure 6. Reduced-period diagram. Reduced period is defined as $P\sqrt{\ell(\ell + 1)}$.

gives us high confidence that those modes have been correctly identified.

The reduced-period spacing between the trapped modes is roughly 2000 s and this observation agrees with the third trapped mode. In the models of Charpinet et al. (2000), this spacing can be taken as a direct measure of the mass of the H envelope, and it was estimated by Charpinet et al. (2000) that it should be of the order of 2000 s, in excellent agreement with our observation. Charpinet et al. (2000) found that for some trapped modes the minimum in period spacing does not correspond to the minimum of the kinetic energy, as expected, but to the maximum instead. Such trapped modes will not be spaced by around 2000 s, but somewhat closer, and will not overlap between modal degrees. We found two spacings, both close

to 2000 s. According to Charpinet et al. (2000), our observations suggest that only the He/H transition zone contributes to the trapping, while the impact from the other, C–O/He, zone is negligible. Having the spacing measured precisely, we could compare it with the models to find the location of the base of the H envelope and its mass. This task will be done as soon as the specific model of this star is available.

7 SUMMARY AND CONCLUSIONS

We have analysed KIC 10001893, a V1093 Her type sdBV star observed with the *Kepler* spacecraft. The data consists of a nearly uninterrupted 1051 d sequence with 1-min cadence. An amplitude spectrum uncovered a rich content of 110 frequencies. We recovered all 27 frequencies of Baran et al. (2011), and while amplitudes vary somewhat, those frequencies are consistent between Q3.3 and this much longer data set. The amplitude spectrum shows both g and p modes, but g modes are strongly favoured, with only six p modes all below 0.1 ppt in amplitude. So, while KIC 10001893 is clearly a hybrid pulsator, it is quite different from the DW Lyn-type hybrid pulsators that show p- and g modes in similar amounts. Hybridity can be an important feature that provides simultaneous constraints both on core and envelope parameters.

We find that KIC 10001893 is quite an enigmatic object with respect to rotationally split multiplets, a feature normally easily seen in similar stars. We found no multiplets among either p or g modes, with the most likely explanation being a very low inclination of the star. This would then be the second sdBV star oriented pole on (Baran et al. 2015b). We identified the modal degree for 50 of the detected g modes, which is about half of the modes, and includes all of the modes with amplitudes higher than 0.1 ppt. Out of those 50 modes, 32 are dipole modes and 18 are quadrupole modes. We calculated échelle diagrams for $l = 1$ and 2 (Fig. 5) and they show wavy excursions from a straight ridge, which are indicative

of non-uniform density profiles in the stellar interior. This feature is similar to ones detected in other *Kepler*-observed sdBV stars (see e.g. Baran & Winans 2012; Østensen et al. 2014; Foster et al. 2015; Baran et al. 2017).

After identifying almost complete sequences of consecutive radial orders for both $l = 1$ and 2, including several trapped modes, we were able to calculate a reduced-period diagram showing almost perfect alignment of the two sequences. The location of the trapped modes, and particularly the spacing between them, provides a useful tool to examine the stellar interior and test various modelling prescriptions against each other. The trapping can be caused by either the transition region of He/H at the base of the H-rich envelope, or deeper trapping regions associated with the transition region between the convective and radiative part of the core. However, the models published thus far (e.g. Charpinet et al. 2000; Ghasemi et al. 2017) do not cover the frequencies in question, neither provide a convincing explanation of the g-mode-trapping structure. Fig. 6 shows three trapped modes, with two of them overlapping between $l = 1$ and 2 sequences. This pattern is almost identical to the one seen in KIC 10553698A, but with all three trapped modes shifted to slightly higher periods. The overlap between the two sequences is so close that in spite of the lack of multiplets as a cross-check for mode identifications, we are confident of our conclusions, at least for the region of overlap. We consider KIC 10001893 to be one of the most prominent sdB stars for applying asteroseismology to infer the physics of the deep interiors of core-helium burning stars.

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